

BISPECTRUM OF THE SUNYAEV-ZEL'DOVICH EFFECT

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ABSTRACT

We perform a detailed study of the bispectrum of the Sunyaev-Zel'dovich effect. Using an analytical model for the pressure profiles of the intracluster medium, we demonstrate the SZ bispectrum to be a sensitive probe of the amplitude of the matter power spectrum parameter σ_8 . We find that the bispectrum amplitude scales as $B_{\text{tSZ}} \propto \sigma_8^{11-12}$, compared to that of the power spectrum, which scales as $A_{\text{tSZ}} \propto \sigma_8^{7-9}$. We show that the SZ bispectrum is principally sourced by massive clusters at redshifts around $z \sim 0.4$, which have been well-studied observationally. This is in contrast to the SZ power spectrum, which receives a significant contribution from less-well understood low-mass and high-redshift groups and clusters. Therefore, the amplitude of the bispectrum at $\ell \sim 3000$ is less sensitive to astrophysical uncertainties than the SZ power spectrum. We show that current high resolution CMB experiments should be able to detect the SZ bispectrum amplitude with high significance, in part due to the low contamination from extra-galactic foregrounds. A combination of the SZ bispectrum and the power spectrum can sharpen the measurements of thermal and kinetic SZ components and help distinguish cosmological and astrophysical information from high-resolution CMB maps.

Subject headings: cosmology: dark matter — galaxies: clusters: general — intergalactic medium

1. INTRODUCTION

The Sunyaev-Zel'dovich (SZ) effect (Sunyaev & Zeldovich 1980) is the principal source of CMB temperature anisotropy power at angular scales smaller than a few arcminutes. It comprises of two components, known as the ‘thermal’ and ‘kinetic’ effects. The former arises from inverse Compton scattering of CMB photons off hot electrons pervading galaxy cluster environments, distorting the Planckian form of the CMB spectrum. The kinetic SZ effect is caused by the Doppler shifting of CMB photons via Thomson scattering off clouds of electrons with a non-zero bulk velocity (along the line-of-sight) relative to the CMB rest frame. Recently, both the South Pole Telescope (SPT, Lueker et al. 2010; Shirokoff et al. 2011; Reichardt et al. 2011) and the Atacama Cosmology Telescope (ACT, Fowler et al. 2010; Dunkley et al. 2011) have placed constraints on the amplitudes of the kinetic and thermal SZ contributions to the small-scale CMB power spectrum. For instance, Reichardt et al. (2011) measured the thermal SZ amplitude at 150 GHz to be $3.69 \pm 0.65 \mu\text{K}^2$ and placed an upper limit on the kSZ amplitude of $< 2.8 \mu\text{K}^2$ (at 95% confidence and assuming no SZ-point source correlation).

Measurements of the thermal SZ power spectrum probe the abundance-weighted pressure profiles of hot gas in groups and clusters over a wide range of redshifts. This is interesting in a cosmological context because the dependence on the abundance of clusters produces a sensitive scaling with σ_8 ; i.e., $C_l \propto \sigma_8^{(7-9)}$ (Komatsu & Seljak 2002; Trac et al. 2011; Shaw et al. 2010, denoted as S10 hereafter). However, efforts to constrain σ_8 in this way have been inhibited by a large

theoretical uncertainty on the amplitude of the tSZ signal (for a fixed cosmological model). Precise predictions of the SZ power spectrum have been hampered by the uncertainties in modeling the state of the intra-cluster medium (ICM) over a broad range of mass and redshift, and at large cluster-centric radii (Battaglia et al. 2010; Trac et al. 2011, S10). Early models and simulations produced predictions for the amplitude of the thermal SZ power spectrum that are discrepant with recent measurements by more than a factor of 2 (Lueker et al. 2010). Recent work, using hydrodynamical simulations (Battaglia et al. 2010, 2011), N-body simulations+semi-analytic gas models (Trac et al. 2011) and analytic models (S10, Efstathiou & Migliaccio 2011; Chaudhuri & Majumdar 2011), have significantly reduced the tension between the observed and predicted values. However the distribution of amplitudes between different models and simulations is still significantly larger than the measurement errors (Reichardt et al. 2011), degrading the constraints that can be placed on cosmological parameters. Furthermore, as current experiments do not cover a sufficiently broad frequency range to enable the tSZ and kSZ signals to be distinguished spectrally, the kSZ contribution must currently be modeled and subtracted from the total SZ power, providing an additional source of systematic uncertainty.

In this work we propose the SZ bispectrum as a new, robust method for constraining the thermal SZ amplitude. The bispectrum, or its real-space counterpart, the 3-point function, is widely used to measure the deviation of a signal from Gaussianity both in the CMB and in large scale structure data. The bispectrum in the CMB map can be measured by using three CMB tem-

perature values and taking the average over the survey sky. The primordial CMB bispectrum can be used to search for signatures of primordial non-Gaussian fluctuations (see e.g., Komatsu et al. 2011, and references therein). The dominant secondary anisotropy signals in the CMB temperature power spectrum arise from lensing of the CMB by large-scale structure, the ISW effect and the SZ effect. These signals, either alone or from their cross correlation, give rise to a measurable bispectrum (Cooray 2000; Komatsu & Spergel 2001; Cooray & Hu 2000). The full non-Gaussian information (primordial or from low-redshift structure) can be obtained by Fourier transforming the CMB temperature field and taking the product of three temperature values that form a triangle, then summing over all possible triangle configurations in harmonic space, assuming all harmonic modes forming the triangles are independent.

One can also measure a skewness spectrum in which two sides of the triangle are summed over to measure the skewness as a function of the third side (Munshi & Heavens 2010). The total skewness spectrum in a CMB map consists of that generated by CMB lensing, the thermal and the kinetic SZ effects, point sources, a signal arising due to the cross correlation of these components, and any primordial contribution. A significant fraction of the kinetic SZ power spectrum is generated by Gaussian density fluctuations in the linear regime (Shaw et al. 2011). Although patchy reionization may contribute a significant fraction of the kSZ spectrum (Mesinger et al. 2012), it is reasonable to expect that the kinetic SZ bispectrum is very small – the overall kSZ amplitude is 50-75% smaller than the tSZ. The SZ skewness spectrum is thus likely to be dominated by the thermal SZ effect. As the primary CMB power decreases rapidly at $\ell > 2000$, the CMB lensing bispectrum should also become negligible at high ℓ . Cooray et al. (2000) have shown that the SZ skewness can be measured using the Planck data with a S/N ≈ 10 , assuming a perfect subtraction of the primary CMB and other foregrounds.

In this work, we focus on the SZ bispectrum, showing that it can provide information complementary to that from the SZ power spectrum. First, as the SZ skewness spectrum is dominated by the thermal SZ component it does not suffer from the observed thermal-kinetic SZ degeneracy. A combination of the SZ skewness and power spectra could therefore help disentangle the thermal and the kinetic signals. We show that the data from current small-scale CMB experiments have the potential to constrain the thermal SZ amplitude using bispectrum measurements with 5 – 10% accuracy. Furthermore, a tightened constraint on the kinetic SZ amplitude can also improve our understanding of the reionization scenario (Shaw et al. 2011; Zahn et al. 2011).

Second, as we will show, the signal is dominated by massive clusters at intermediate redshift for which high-precision X-ray observations exist. This is in contrast to the power spectrum where the signal mainly comes from the lower mass and higher redshift groups and clusters (e.g., Trac et al. 2011). The theoretical uncertainty in the SZ skewness spectrum is thus expected to be significantly smaller than that of the SZ power spectrum. Combined measurements of the power spectrum and the bispectrum can thus be used to distinguish the contribution to the power spectrum from different cluster mass

and redshift ranges.

The goal of this study is to derive the SZ skewness spectrum using an analytic model for the ICM gas and analyze its astrophysics and cosmology dependence. In Section 2 we discuss the theory of the SZ bispectrum, and in Section 3 we describe our model for calculating the thermal pressure of the intra-cluster medium used in the bispectrum calculation. In Section 4 we present our results, including the distribution of the skewness spectrum signal over mass and redshift, the theoretical uncertainty, cosmological dependence, and the prospects of detectability from the current data from SPT like surveys. Section 5 presents summary and discussions. Throughout this work, we assume a fiducial cosmology with zero curvature and parameters: $\Omega_b = 0.045$, $\Omega_m = 0.27$, $\sigma_8 = 0.8$, $h = 0.71$, $n_s = 0.97$, (Komatsu et al. 2011) and $w_0 = -1$.

2. THEORETICAL FRAMEWORK

2.1. Sunyaev-Zel'dovich effect

The Compton y -parameter for the thermal SZ effect is

$$y = \frac{\Delta T}{T_{\text{CMB}}} = \left(\frac{k_B \sigma_T}{m_e c^2} \right) \int n_e(l) T_e(l) dl, \quad (1)$$

where l is the distance along the line-of-sight; n_e and T_e are the electron number density and temperature, respectively.

The thermal SZ power spectrum can be calculated by summing up the square of Fourier transform of the projected SZ profile, weighted by the number density of clusters of a given mass and redshift,

$$C_l = f(x_\nu)^2 \int dz \frac{dV}{dz} \int d \ln M \frac{dn(M, z)}{d \ln M} \tilde{y}^2(M, z, \ell), \quad (2)$$

where $V(z)$ is the comoving volume per steradian, $f(x_\nu) = x_\nu (\coth(x_\nu/2) - 4)$ with $x_\nu = h\nu/(k_B T_{\text{CMB}})$ at frequency ν and is negative (positive) below (above) 220 GHz, T_{CMB} is the temperature of CMB at $z = 0$ and $n(M, z)$ is the halo mass function of Tinker et al. (2008), $\tilde{y}(M, z, \ell)$ is the Fourier transform of the projected SZ profile of a cluster, given by

$$\tilde{y}(M, z, \ell) = \frac{4\pi r_s}{\ell_s^2} \left(\frac{\sigma_T}{m_e c^2} \right) \int_0^\infty dx x^2 P_e(M, z, x) \frac{\sin(\ell x/\ell_s)}{\ell x/\ell_s}, \quad (3)$$

where $x = r/r_s$, $\ell_s = D_A(z)/r_s$, r_s is the scale radius of the 3D pressure profile, $D_A(z)$ is the angular diameter distance at redshift z and P_e is the electron pressure profile. Note that we use the mass function fit for 400 times the *mean* density, whereas the pressure profiles are defined at the *virial* overdensity (Bryan & Norman 1998). We convert the mass function definition to the *virial* density in Eq. 2. The cosmology dependence of the SZ power spectrum comes from the angular diameter distance and the growth function while the astrophysical dependence comes from the pressure profiles.

2.2. Bispectrum Theory

The CMB temperature fluctuations $\Delta\theta(\hat{\mathbf{n}})$ in a certain direction, $\hat{\mathbf{n}}$, can be expanded into spherical harmonics as

$$a_{\ell m} = \int d^2 \hat{\mathbf{n}} \frac{\Delta T}{T} Y_{\ell m}^*(\hat{\mathbf{n}}). \quad (4)$$

The angular bispectrum is

$$B_{\ell_1 \ell_2 \ell_3}^{m_1 m_2 m_3} = \langle a_{\ell_1 m_1} a_{\ell_2 m_2} a_{\ell_3 m_3} \rangle, \quad (5)$$

where the angle-averaged quantity in the full sky limit can be written as

$$B(\ell_1 \ell_2 \ell_3) = \sum_{m_1 m_2 m_3} \begin{pmatrix} \ell_1 & \ell_2 & \ell_3 \\ m_1 & m_2 & m_3 \end{pmatrix} B_{\ell_1 \ell_2 \ell_3}^{m_1 m_2 m_3}, \quad (6)$$

which has to satisfy the conditions: $m_1 + m_2 + m_3 = 0$, $\ell_1 + \ell_2 + \ell_3 = \text{even}$, and $|\ell_i - \ell_j| \leq \ell_k \leq \ell_i + \ell_j$. Here $B(\ell_1 \ell_2 \ell_3)$ is the bispectrum in the full-sky limit. In the small angle limit, the flat-sky approximation is valid. By defining $b(\ell_1 \ell_2 \ell_3)$ to be the bispectrum in the flat-sky limit, then the correspondence between the full-sky and the flat-sky bispectrum is (Hu 2000),

$$B(\ell_1 \ell_2 \ell_3) \approx \sqrt{\frac{(2\ell_1 + 1)(2\ell_2 + 1)(2\ell_3 + 1)}{4\pi}} \begin{pmatrix} \ell_1 & \ell_2 & \ell_3 \\ 0 & 0 & 0 \end{pmatrix} 7 \times b(\ell_1 \ell_2 \ell_3).$$

Throughout this work, we focus on the high- ℓ regime (to avoid confusion from primary CMB and its lensing), where a flat sky approximation is valid. We will refer to the flat-sky bispectrum; namely, $b(\ell_1 \ell_2 \ell_3)$ as the bispectrum for brevity and $B(\ell_1 \ell_2 \ell_3)$ as the full-sky bispectrum. Note that, for $\ell > 500$, the Wigner 3-j symbol can be approximated as

$$\begin{pmatrix} \ell_1 & \ell_2 & \ell_3 \\ 0 & 0 & 0 \end{pmatrix} \approx \sqrt{\frac{2}{\pi}} \frac{(-1)^{L/2}}{[L(L - 2\ell_1)(L - 2\ell_2)(L - 2\ell_3)]^{1/4}} \quad (8)$$

if $L = \ell_1 + \ell_2 + \ell_3$ is even, and vanishes for odd L . The approximate expression for Eq. 8 is valid to better than 1 percent for $\ell > 500$, the ℓ range we are interested in this study.

The thermal SZ bispectrum is the volume integral of the cube of the Fourier transform of the pressure profile weighted by the halo mass function,

$$b(\ell_1 \ell_2 \ell_3) = f(x_\nu)^3 \int dz \frac{dV}{dz} \int d \ln M \frac{dn(M, z)}{d \ln M} \times \tilde{y}(M, z, \ell_1) \tilde{y}(M, z, \ell_2) \tilde{y}(M, z, \ell_3). \quad (9)$$

We sum over the two smaller values of ℓ s and writing the bispectrum as a function of the largest ℓ , we define the skewness spectrum as

$$\Lambda(\ell) = \sqrt{\sum_{\ell_1 \ell_2} b^2(\ell \ell_1 \ell_2)}, \quad (10)$$

where $\ell + \ell_1 + \ell_2$ is even and $\ell_1 \leq \ell_2 \leq \ell$.

In previous work, the amplitude of the thermal SZ power spectrum has been defined in terms of the expected amplitude at $\ell = 3000$ for a given set of cosmological parameters (Lueker et al. 2010; Shirokoff et al. 2011; Reichardt et al. 2011; Dunkley et al. 2011). In this study, we define A_{tSZ} as the ratio of the thermal SZ power spectrum amplitude at non-fiducial cosmology to that at $\sigma_8 = 0.8$ as

$$A_{\text{tSZ}} \equiv \frac{C_{3000}(\sigma_8)}{C_{3000}(\sigma_8 = 0.8)}. \quad (11)$$

Parameters	Fiducial	Range
Dark Matter concentration (A_C)	1	0.8 – 1.2
Energy Feedback (ϵ_f)	4×10^{-7}	$(1-100) \times 10^{-7}$
Dynamical Friction Heating (ϵ_{DM})	0.05	0.0 – 0.1
Non-thermal pressure normalization (α_0)	0.2	0.0 – 0.3
Non-thermal pressure evolution (β)	0.5	-1 – +1

TABLE 1
FIDUCIAL VALUES AND THE UNCERTAINTY RANGE IN THE ICM PARAMETERS.

Similarly, we define the amplitude of the skewness spectrum as

$$B_{\text{tSZ}} \equiv \frac{\Lambda_{3000}(\sigma_8)}{\Lambda_{3000}(\sigma_8 = 0.8)}, \quad (12)$$

where Λ_{3000} is the amplitude of the skewness spectrum at $\ell = 3000$. For the fiducial cosmology, $B_{\text{tSZ}} = A_{\text{tSZ}} = 1$. Given that the skewness spectrum is proportional to the cube of the pressure profile the expected relation between B_{tSZ} and A_{tSZ} is $B_{\text{tSZ}} \propto A_{\text{tSZ}}^{1.5}$, assuming similar mass and redshift distribution of the power spectrum and the bispectrum. In Section 4.3, we study the relation between the SZ amplitude of the power spectrum (A_{tSZ}) and the skewness spectrum (B_{tSZ}) in more detail.

3. MODELING THE INTRA-CLUSTER MEDIUM

To calculate the thermal pressure profiles of groups and clusters, we use the physically motivated analytic model presented in S10. Here we briefly describe the key features of the model, and refer the readers to S10 for more details.

The model assumes that intra-cluster gas resides in hydrostatic equilibrium within the gravitational potential of the host dark matter halo,

$$\frac{dP_{\text{tot}}(r)}{dr} = -\rho_g(r) \frac{d\Phi(r)}{dr}, \quad (13)$$

where $\rho_g(r)$ is the gas density at radius r from the cluster center, the total pressure, $P_{\text{tot}}(r) = P_{\text{th}}(r) + P_{\text{nt}}(r)$, is given by a sum of thermal and non-thermal pressures, and $\phi(r)$ is the gravitational potential. The gas is assumed to have a polytropic equation of state, $P_{\text{tot}} = P_0(\rho_g/\rho_0)^\Gamma$, where $\Gamma = 1.2$ and P_0 and ρ_0 are the central gas pressure and density respectively. The total mass distribution is assumed to follow the Navarro-Frenk-White (NFW) profile (Navarro et al. 1996, 1997),

$$\rho_{\text{tot}}(r) = \frac{\rho_s}{x(1+x)^2}, \quad (14)$$

where ρ_s is the normalization constant, $x = r/r_s$, and r_s is the characteristic radius and can be defined in terms of the concentration of halos, $r_s = r_{\text{vir}}/c_{\text{vir}}$ where “vir” refers to the virial overdensity (Bryan & Norman 1998) with respect to the *critical density* of the universe. We adopt the concentration-mass relation given in Duffy et al. (2008):

$$c_{\text{vir}}(M, z) = 7.85 A_c \left(\frac{M_{\text{vir}}}{2 \times 10^{12} h^{-1} M_\odot} \right)^{-0.081} (1+z)^{-0.71}, \quad (15)$$

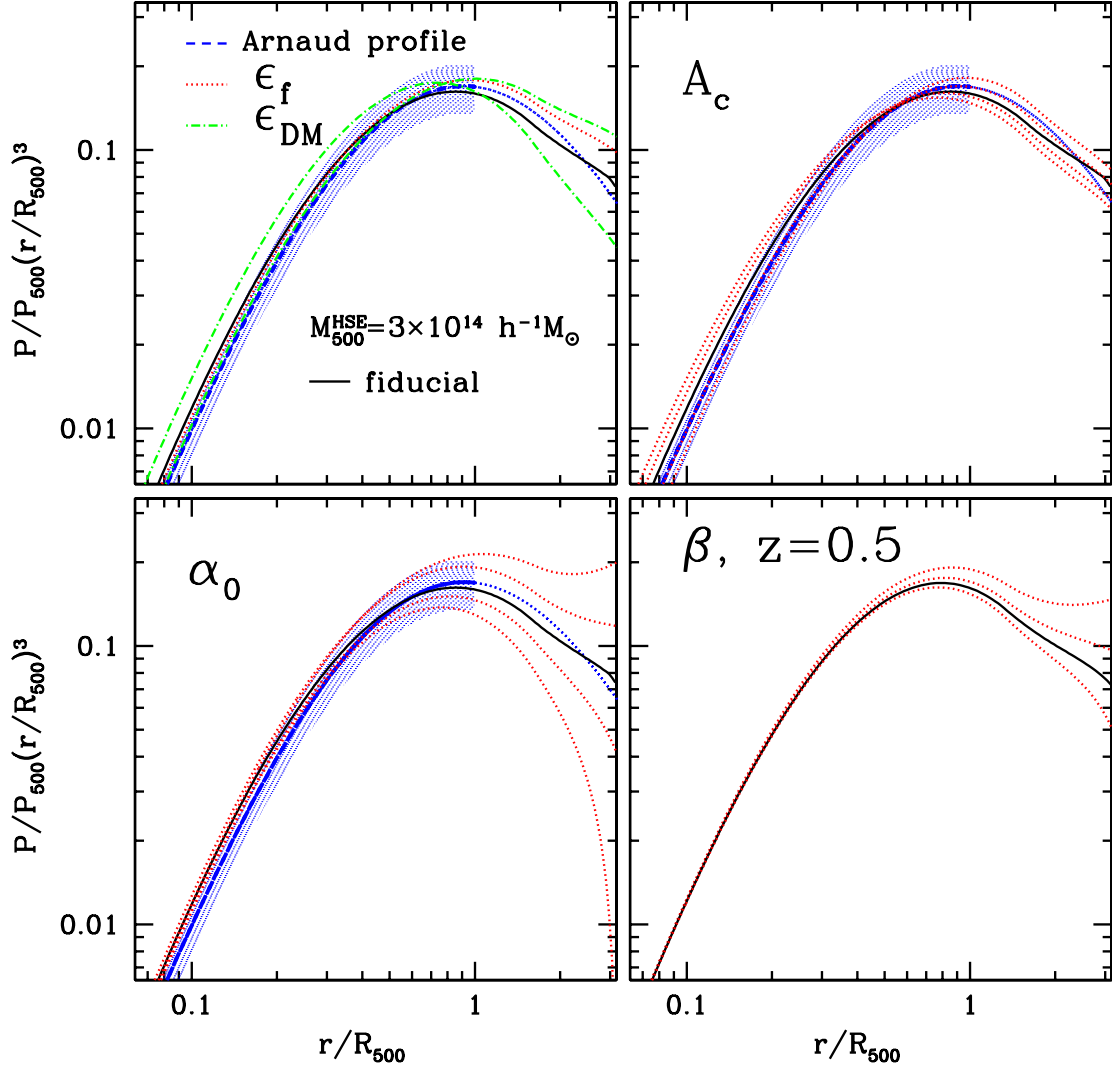


FIG. 1.— Comparison between the observed ‘universal’ pressure profile (scaled with respect to P_{500}) of (Arnaud et al. 2010) (blue dashed line, shaded region denotes 20% scatter) and our model predictions at $z=0$ for a cluster of hydrostatic mass $M_{500}^{HSE} = 3 \times 10^{14} h^{-1} M_{\odot}$. In each panel we vary one of the input model parameters to demonstrate their effect on the pressure profile. The black solid line represent the fiducial values as given in Table 1. The top-left panel shows the profiles obtained when ϵ_f is 10^{-5} (red dotted line) and ϵ_{DM} varies from 0.2 to 0.0 (green dashed lines, from top to bottom at $2R_{500}$); in the top right panel we vary A_c from 0.8, 0.9, 1.1, 1.2 (red dotted lines, top to bottom at $2R_{500}$), the bottom-left panel shows the variation with α_0 from 0.3, 0.24, 0.14, 0.0 (red dotted lines, bottom to top at $2R_{500}$) and the bottom-right panel shows the variation with β at $z = 0.5$ over $\{1, 0.5, 0, -1\}$ where self-similar evolution is factored out (red dotted lines, from bottom to top). Note that the observed result is at $< R_{500}$, the blue dotted line is an extension of the Arnaud profile based on the simulation result.

where A_c is the normalization factor, with the fiducial value $A_c = 1$. Note that the gas is distributed as a massless tracer of dark matter distribution.

The model accounts for non-gravitational processes that affect the thermal properties of the ICM including star formation, energy feedback from supernovae (SNe) and Active Galactic Nuclei (AGN), and non-thermal pressure support due to bulk gas motions and turbulence. Star-formation is implemented by converting a certain (mass and redshift-dependent) fraction of the gas mass into stars. We adopt the observed stellar mass fraction for the local X-ray groups and clusters (Giodini et al. 2009) and assume that the stellar populations evolve according to the ‘fossil’ model of Nagamine et al. (2006).

Following Ostriker et al. (2005), a fraction of the rest mass energy of the stars, $\epsilon_f M_* c^2$, is put back into the gas due to energy feedback from SNe and AGN, where ϵ_f is a free parameter. The model also accounts for the transfer of energy from dark matter to gas by dynamical heating by infalling substructures, determined by the product of the total binding energy of the halo $|E_{DM}|$ and a free parameter ϵ_{DM} . Finally, we take into account the non-thermal pressure support due to random gas motions as seen in hydrodynamical simulations (e.g., Kay et al. 2004; Rasia et al. 2004; Lau et al. 2009; Battaglia et al. 2010; Nelson et al. 2011). We assume the non-thermal pressure to be a certain fraction of the total

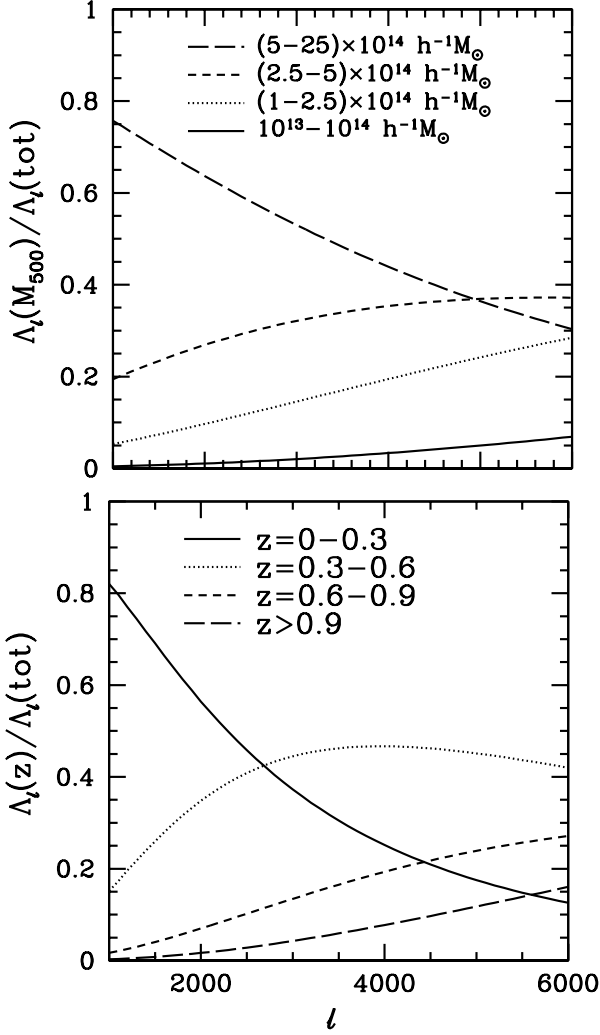


FIG. 2.— Contribution to the SZ skewness spectrum of objects within a given mass range (top) and redshift range (bottom) range. We show the ratio of the skewness spectrum in a mass/redshift range to the total skewness spectrum.

pressure and a power law in radius,

$$\frac{P_{\text{nt}}}{P_{\text{tot}}}(z) = \alpha(z) \left(\frac{r}{R_{500}} \right)^{n_{\text{nt}}}, \quad (16)$$

where $\alpha(z) = \alpha_0 f(z)$, α_0 is the ratio of non-thermal to total pressure at R_{500} , the radius at which the spherical overdensity of the cluster is 500 times the critical density of the universe (enclosing mass M_{500}). $f(z) = \min[(1+z)^\beta, (4^{-n_{\text{nt}}}/\alpha_0 - 1) \tanh(\beta z) + 1]$, n_{nt} and β are the radial and redshift dependence of the non-thermal pressure support, respectively. We set $n_{\text{nt}} = 0.8$, motivated by comparisons to hydrodynamical simulations (S10, Battaglia et al. 2011), leaving β as a free parameter.

In summary, our ICM model has five free parameters A_C , ϵ_f , ϵ_{DM} , α_0 , and β . Since the key observable for the SZ effect is the pressure profile, we use measurements of the pressure profiles by Arnaud et al. (2010) based on 31 massive ($M_{500} > 10^{14} h^{-1} M_\odot$), low redshift ($z \lesssim 0.2$) clusters. Arnaud et al. (2010) showed that the pressure

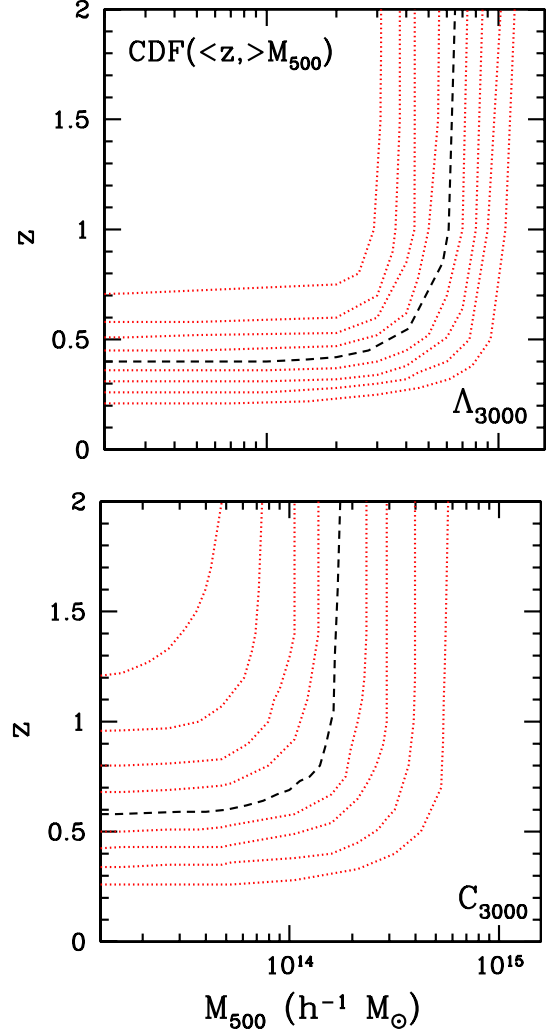


FIG. 3.— Top panel: the distribution of the skewness spectrum over mass and redshift at $\ell = 3000$. At $M_{500} = 5 \times 10^{13} h^{-1} M_\odot$, the lines show the cumulative distribution function, $\text{CDF}(>M, <z)=0.1$ to 0.9 from bottom to top with a step of 0.1 . Bottom panel: the distribution of the SZ power spectrum at $\ell = 3000$ (for comparison with the skewness distribution). The lines represent similar quantities as that in the top panel.

profiles in their sample adheres to a universal form with $\sim 20\%$ scatter. Figure 1 shows the pressure profile fit to the Arnaud et al. (2010) data along with our predictions for a typical cluster of mass $M_{500}^{\text{HSE}} = 3 \times 10^{14} h^{-1} M_\odot$ at $z=0$, where M_{500}^{HSE} is a cluster mass estimated assuming the hydrostatic equilibrium (see S10 for more discussions). Each of the 5 parameters are varied over a wide range, including some extreme cases with a very high feedback parameter or zero non-thermal pressure.

The parameter ϵ_f determines the amount of non-gravitational (‘feedback’) energy injected into the ICM. High levels of feedback inflates the gas distribution, lowering the pressure in the inner region while boosting it in the cluster outskirts. For cluster-mass halos, the feedback energy is a small fraction of the binding energy and has a minimal effect on the pressure profile. The upper-left panel of Figure 1 shows how ϵ_f changes the pressure profile in a cluster-mass halo at $z = 0$.

Increasing ϵ_f by factor of ~ 100 changes the pressure profile by few percent at radii $R < R_{500}$ and increases to about 30% at $3R_{500}$. This is in qualitative agreement with the findings from hydrodynamical simulations including AGN feedback which indicate the impact of AGN feedback is more pronounced in groups than in clusters (Bhattacharya et al. 2008; Puchwein et al. 2008; Sijacki et al. 2007; Teyssier et al. 2011; McCarthy et al. 2010; Battaglia et al. 2010, 2011; Short et al. 2012). We note however that our feedback model is highly simplistic and does not capture the full detail of AGN feedback as implemented in hydrodynamical simulations (for example, energy injection via jets).

ϵ_{DM} determines the fraction of the total dark matter energy transferred to the ICM during mergers. Clusters frequently undergo mergers and hence ϵ_{DM} may have a substantial effect on the ICM physics of clusters. As shown in Figure 1 (top-left), ϵ_{DM} changes the pressure profile by 20-30% over the entire radius range as ϵ_{DM} varies from 0 – 0.2.

The normalization of the concentration-mass relation (A_c) determines the characteristic radius of an NFW dark matter profile, approximately corresponding to the radius at which the slope changes from -1 to -3 . The upper-right panel of Figure 1 shows the change in the thermal pressure profile as A_c varies from 0.8 – 1.2. Increasing the concentration deepens the central potential which draws the gas towards the cluster core, steepening the pressure profile. The impact of increasing halo concentration provides a similar (although opposite) effect to that of energy feedback.

The nonthermal pressure parameter, α_0 , determines the fraction of the total gas pressure contributed by random gas motions. As α_0 is varied is the thermal pressure fraction adjusts accordingly such that the total gas pressure is fixed. As shown in Figure 1 (bottom-left), the thermal pressure profile changes by about 60% at R_{500} as α_0 varies from 0 – 0.3. β determines the redshift evolution of the non-thermal pressure fraction. S10 obtained an estimate for β by comparing the redshift 0 and 1 outputs of the simulations of Lau et al. (2009), finding that the best fit value for α_0 increases from 0.18 – 0.26 implying $\beta = 0.5$. The bottom-right panel in Figure 1 shows the pressure profile of a cluster at $z=0.5$ and how β changes the thermal pressure profile. A positive value of β increases the amount of non-thermal pressure at higher redshift and hence decreases the thermal pressure as redshift increases. The pressure profile decreases by about 20% as β varies from -1 to 1.

Overall, we find that the range $\epsilon_f < 10^{-6}$, $\epsilon_{DM} = 0.04 - 0.06$, $A_c = 0.8 - 1.1$, and $\alpha_0 = 0.14 - 0.24$ to be in agreement with the Arnaud pressure profile within the uncertainty (see Figure 1). Note that, as described in S10, the constraints on $\epsilon_f < 10^{-6}$ and $\epsilon_{DM} = 0.04 - 0.06$ mainly come from the gas fraction-mass relation of Vikhlinin et al. (2006). We also find the entropy scaling relation (e.g., entropy - X-ray temperature relation) predicted by our ICM model to be in agreement with the data of Pratt et al. (2010) within the ranges of the ICM parameters. This illustrates that the current observations can indeed constrain the ICM parameter space significantly.

4. THE SZ SKEWNESS SPECTRUM

Mass Range $M_{500}[h^{-1}M_{\odot}]$	power %	z-range	power %
$< 10^{14}$	3	0–0.3	38
$(1 - 2.5) \times 10^{14}$	15	0.3–0.6	45
$(2.5 - 5) \times 10^{14}$	30	0.6–0.9	12
$(5 - 25) \times 10^{14}$	52	> 0.9	5

TABLE 2
DISTRIBUTION OF THE SKEWNESS SPECTRUM OVER DIFFERENT MASS RANGE AT $\ell \sim 3000$.

4.1. Mass and redshift contributions

Figure 2 shows the fraction of skewness spectrum signal contributed by different ranges of halo mass and redshift. At $\ell \sim 3000$, galaxy groups ($M_{500} < 10^{14}h^{-1}M_{\odot}$) source only 5% of the signal, while lower mass clusters contribute about 20%. The majority of the signal comes from the massive clusters. The distribution of the skewness spectrum is tabulated in Table 2 over different mass and redshift ranges. We also calculate the signal expected from clusters that are below the detection threshold of the current generation of SZ surveys. Following Benson et al. (2011); Vanderlinde et al. (2010), we choose the SPT mass threshold of $M_{500} \sim 2.5 \times 10^{14}h^{-1}M_{\odot}$, roughly corresponding to the detection significance of 5σ . The skewness spectrum gets only 20% (at $\ell = 1500$) to 40% (at $\ell = 6000$) of the signal from clusters with $M_{500} \lesssim 2.5 \times 10^{14}h^{-1}M_{\odot}$. In contrast, the SZ power spectrum gets about two-thirds of the signal from groups and clusters with $M_{500} \lesssim 2.5 \times 10^{14}h^{-1}M_{\odot}$.

Figure 3, top panel shows the cumulative fractional distribution of the skewness spectrum for halos with mass $< M_{500}$ and redshift $> z$, from 10-90% of the signal (top-left to bottom-right) at $\ell = 3000$. In terms of the redshift distribution, about 90% of the skewness spectrum signal comes from $z < 1$. The bottom panel of Figure 3 shows the distribution of the SZ power spectrum, showing that a significant fraction of the thermal SZ power comes from high- z , low-mass objects (i.e., about 40% of the power coming from $z > 0.7$ and $M_{500} < 10^{14}h^{-1}M_{\odot}$) where the ICM properties are still very uncertain. In comparison, the SZ skewness spectrum signal is dominated by low- z , massive clusters.

4.2. Astrophysical Uncertainties

We now evaluate how robust the predicted skewness spectrum is to uncertainties in the underlying ICM model parameters. Figure 4 shows the variation of the SZ skewness spectrum over the range of ICM parameters given in Table 1 and plotted in Figure 1. As discussed in Section 3, ϵ_f does not have a significant effect on the massive clusters that dominate the skewness spectrum. Consequently, as shown in the top-left panel of Figure 1, the skewness spectrum varies only by 14% as ϵ_f varies from $(1-100) \times 10^{-7}$. On the other hand, the pressure profiles of massive clusters vary substantially with ϵ_{DM} , and the skewness spectrum varies by about 40% as ϵ_{DM} is varied from 0 to 0.2. In contrast, the power spectrum receives a significant contribution from a wider range in mass, $M_{500} = 10^{13} - 10^{15}h^{-1}M_{\odot}$, especially at high ℓ . As a result, ϵ_f shows more variations at $\ell = 8000$ (group scales) than at $\ell = 2000$ (cluster scales, see Fig. 7 of S10).

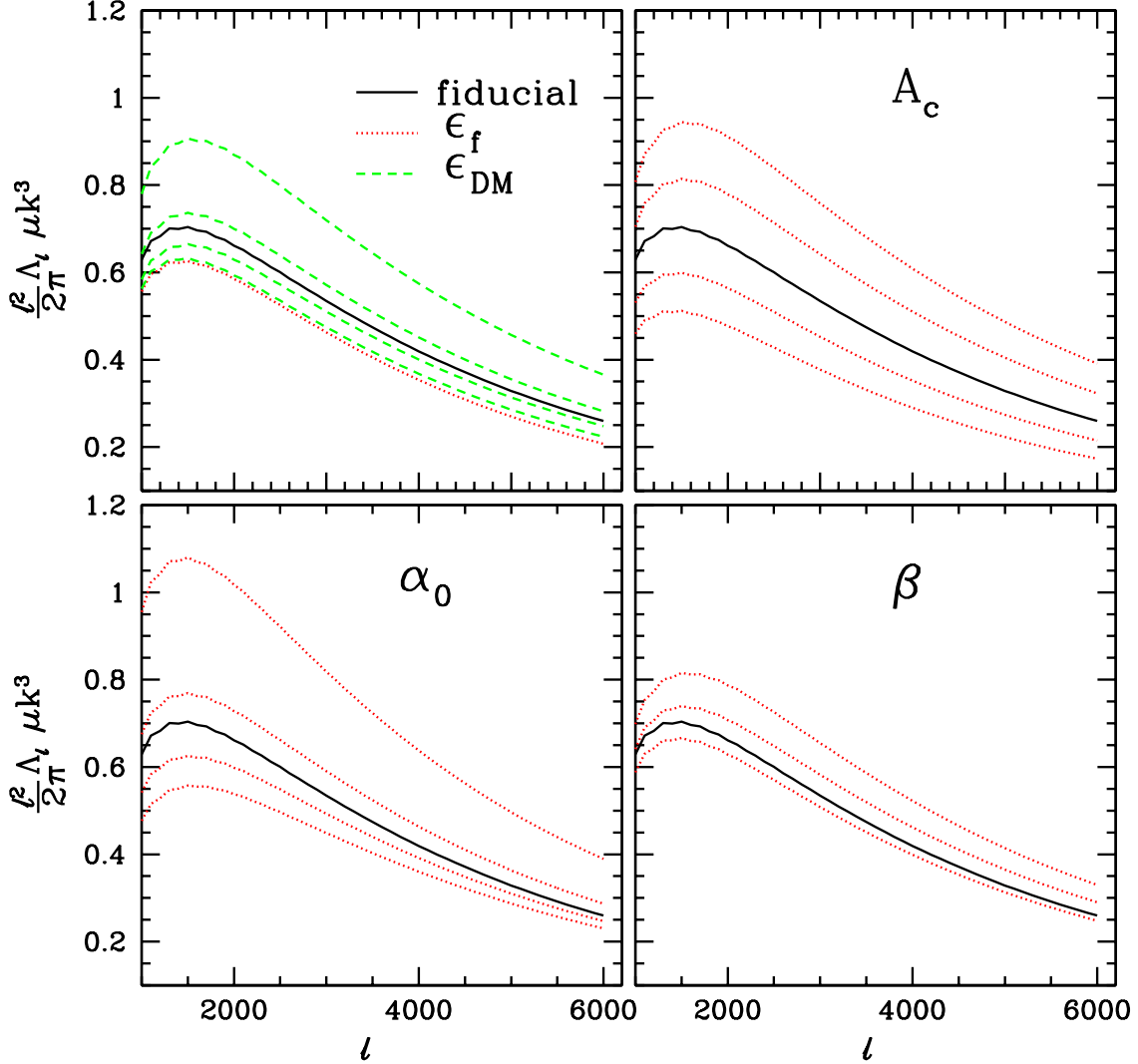


FIG. 4.— Dependence of the SZ skewness spectrum on the cluster gas physics. The solid line indicates the fiducial model with $\epsilon_f = 4 \times 10^{-7}$, $\epsilon_{DM} = 0.05$, $A_c = 1.0$, $\alpha_0 = 0.2$, and $\beta = 0.5$. The dotted and dashed lines represent the skewness spectrum predicted by our ICM model for a range of ICM parameter values. The top-left panel shows the impact of ϵ_f (red dotted line: 10^{-5} and black solid line: 4×10^{-7}) and ϵ_{DM} (green dashed lines: 0.0, 0.04, 0.06, and 0.2 from top to bottom); the top-right panel shows the change with A_c (0.8, 0.9, 1.0, 1.1, 1.2 from bottom to top); the bottom-left panel shows the variations with α_0 (0.0, 0.14, 0.2, 0.24, 0.3 from top to bottom); and the bottom-right panel shows the variation with β (−1, 0, 0.5, 1 from top to bottom).

Increasing the normalization of the concentration-mass relation (A_c) deepens the potential well of the clusters which steepens the pressure profiles. The top-right panel of Figure 4 shows the change in the skewness spectrum is about $\pm 30\%$ for a $\pm 20\%$ change in A_c (compared to the power spectrum where the variation is $\sim 20\%$). This can be understood from the fact that the skewness amplitude is $\propto (\text{power spectrum})^{1.5}$, so a 20% change in power spectrum corresponds to a 30% change in the skewness spectrum. Also note that in contrast to the case of ϵ_f , the effect of A_c is not significantly mass-dependent.

The lower-left panel shows the effect on the skewness spectrum of varying the non-thermal pressure parameter α_0 . The skewness spectrum changes by about 70% across the range $\alpha_0 = 0-0.3$. As the SZ skewness spectrum gets most of its signal from massive clusters at $z \lesssim 0.4$, the

uncertainty in the z -evolution of the non-thermal pressure (β) contributes only about 20% uncertainty to the SZ skewness spectrum. The SZ power spectrum, on the other hand gets its signal from high- z objects and as a result β adds $\sim 30-40\%$ uncertainty in the SZ power spectrum (bottom-right panel of Figure 7 in S10).

Figure 5 shows the astrophysical uncertainties in the shaded area around the fiducial model. Here, the total uncertainty is computed by adding the uncertainties (calibrated by the observed pressure profile and discussed in detail in Section 3) in each ICM parameter in quadrature. On average, about 75–80% of the skewness spectrum signal arises from massive lower redshift objects; i.e., $M_{500} \gtrsim 5 \times 10^{14} h^{-1} M_\odot$ at $\langle z \rangle 0.4$, and hence is less sensitive to the variation in ICM parameters than the power spectrum (c.f. Figure 7 in S10). Figure 5 shows

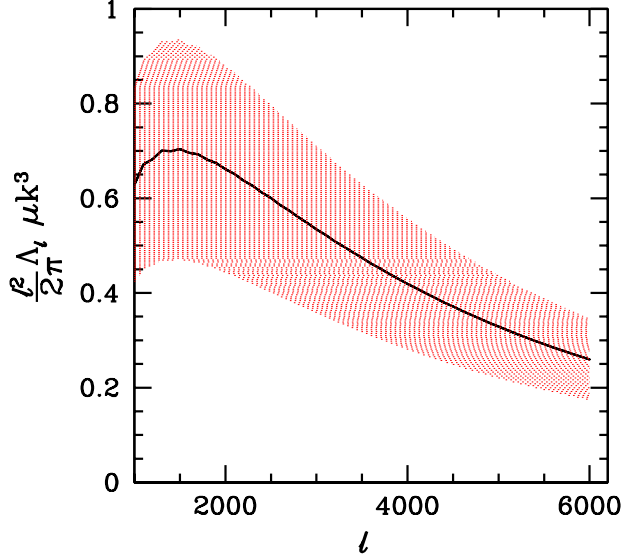


FIG. 5.— Astrophysical uncertainties in the SZ skewness spectrum. The solid line represents the fiducial model and the shaded region indicates astrophysical uncertainty for the ICM parameter range calibrated by the observed pressure profile (see Section 3). At $\ell = 3000$, the astrophysical uncertainties are $\sim 33\%$.

the gas physics uncertainty in the skewness spectrum to be about 33%. We also provide the theoretical prediction of the skewness spectrum for the fiducial values of cosmological and ICM parameters.

At this point, it is worth commenting on effects of several simplifying assumptions in our model. First, we assumed a constant $\Gamma = 1.2$. However, recent work suggests that Γ could be radially dependent, especially at large radii ($r < R_{\text{vir}}$). However, the bulk ($\approx 95\%$) of the skewness signal arises from $r < R_{\text{vir}}$, and $\Gamma \approx 1.14 - 1.3$ over the redshift range $z = 0 - 1$ (Battaglia et al. 2011). We therefore find that the skewness spectrum changes by about $\pm 10 - 15\%$ in the specified range of Γ . Also, there is uncertainty in the number of baryons converted into stars in clusters, since this is only constrained by redshift zero observations. Secondly, the redshift dependence of the model assumes that evolution of the stellar fraction is the same for all galaxies, although it is possible that the star fraction evolve differently in a cluster environment. In order to assess the impact of star formation model on the skewness spectrum, we considered a no evolution model and found that the difference between the “no-evolution” model and the fiducial model (namely the “fossil” model of Nagamine et al. (2006)) is negligible. Finally, recent work on the thermal SZ power spectrum have shown that variations about a mean profile may produce up to a 15% excess in the spectrum at $\ell \sim 3000$ (e.g., Battaglia et al. 2011). This variation is non-Gaussian and radially-dependent and can lead to the enhancement of the bispectrum amplitude.

4.3. The relation between the power spectrum and the skewness spectrum amplitude

Given a measurement of the SZ skewness spectrum amplitude, we can predict the expected amplitude of the thermal SZ power spectrum, namely the $B_{\text{tSZ}} - A_{\text{tSZ}}$

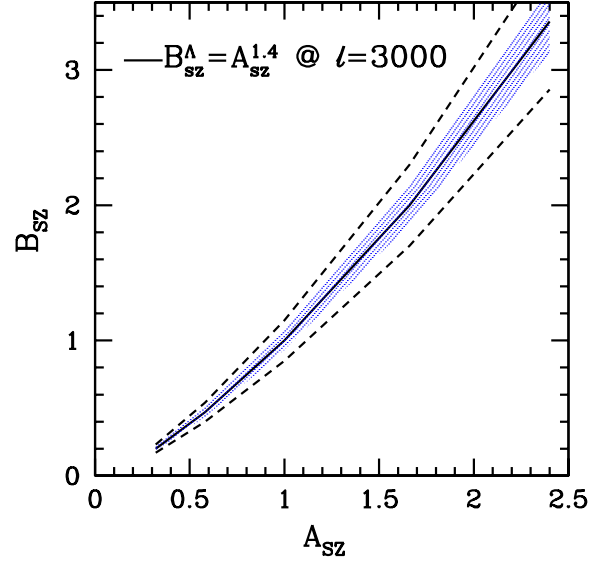


FIG. 6.— Scaling of $B_{\text{tSZ}} - A_{\text{tSZ}}$ for the skewness spectrum when σ_8 varies from 0.7 to 0.9. The relations are relatively robust to the changes in gas physics and vary by only $\sim 7\%$ (shaded area) for the allowed ICM parameter range calibrated by the observed pressure profile (see Section 3). The dashed lines indicate the uncertainty range bracketed by the two extreme gas physics scenarios.

ℓ	$\ell^2 / (2\pi) \Lambda(\ell) \mu K^3$
1000	0.66
2000	0.69
3000	0.56
4000	0.44
5000	0.35
6000	0.28

TABLE 3
THE THEORETICAL PREDICTION OF THE SKEWNESS SPECTRUM FOR THE FIDUCIAL VALUES OF ICM AND COSMOLOGY PARAMETERS.

relation. Recall that the power spectrum and the skewness spectrum are proportional to the square and the cube of the pressure profile, respectively. If we assume the skewness spectrum and the power spectrum have the same mass and redshift distribution, then any changes in gas physics affects the skewness spectrum and (power spectrum)^{1.5} in the same way. However, as shown in Sec 4.1, the skewness spectrum depends mostly on the massive clusters at intermediate redshift, while the power spectrum receives a substantial contribution from higher redshift galaxy groups. Furthermore, the cluster physics affect the gas properties in lower mass structures more than in clusters. Therefore, the $B_{\text{tSZ}} - A_{\text{tSZ}}$ relation is expected to deviate from the above relation and depends on both cosmology and cluster astrophysics.

To investigate the dependence of the $B_{\text{tSZ}} - A_{\text{tSZ}}$ relation on cluster physics and cosmology, we vary σ_8 from 0.7 – 0.9 in steps of 0.05 and compute the $B_{\text{tSZ}} - A_{\text{tSZ}}$ relation for each value of σ_8 for our fiducial gas physics model. Figure 6 shows that the skewness spectrum amplitude is proportional to the power spectrum amplitude,

$$B_{\text{tSZ}}(\ell = 3000) = A_{\text{tSZ}}^{1.4}(\ell = 3000). \quad (17)$$

This is shallower than the expected $A_{\text{tSZ}}^{1.5}$ behavior, primarily because the skewness spectrum signal comes from

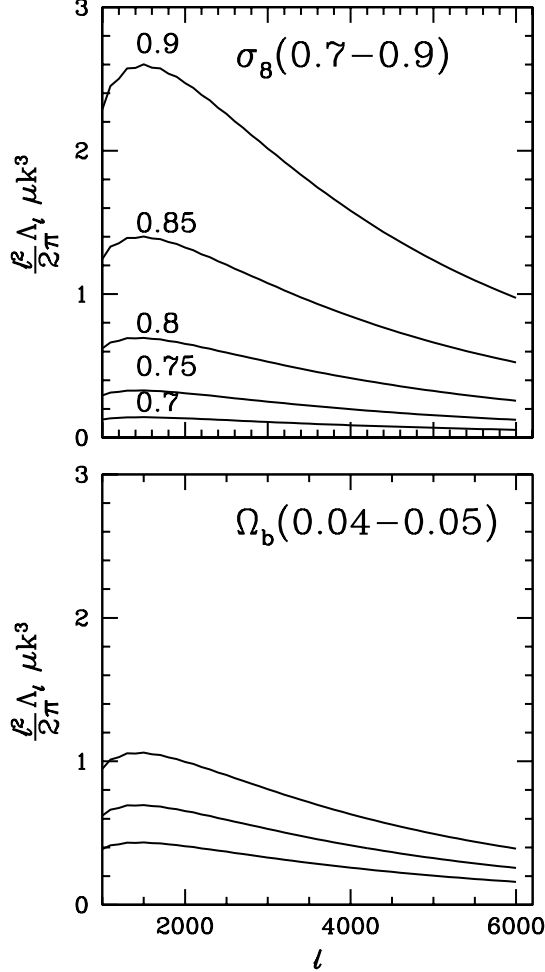


FIG. 7.— Dependence of the skewness spectrum on σ_8 and Ω_b . The fiducial values are $\Omega_b = 0.045$ and $\sigma_8 = 0.8$. The range of parameters are indicated in the top-right corner of each panel. The parameter values increase from bottom to top.

more massive clusters which are not point-like at the angular scales equivalent to $\ell \sim 3000$. We thus study the radial dependence of the skewness spectrum signal and find that less than 5% signal comes from $[1-2]R_{\text{vir}}$, 20% comes from $[0.5-1]R_{\text{vir}}$, 60% from $[0.2-0.5]R_{\text{vir}}$, the rest 15% from $\leq 0.2R_{\text{vir}}$. Thus a simple scaling of the pressure normalization with mass cannot explain the $B_{\text{tSZ}} - A_{\text{tSZ}}$ relation. Next we investigate the variation of the scaling relation with the ICM parameters. Within the range of the ICM parameters allowed by the observed pressure profile, we find the $B_{\text{tSZ}} - A_{\text{tSZ}}$ relation is robust and the overall uncertainty does not vary by more than 7% (indicated by shaded area in Figure 6). The tight scaling arises because both the power spectrum and the skewness spectrum depend on the pressure profile. As a result, the amplitudes of the two spectra varies in similar ways as we vary our model parameters.

To check the robustness of the relation further, we consider two extreme gas physics scenarios, such that the SZ power spectrum varies from $2.5 - 10 \mu\text{K}^2$ (roughly by factor of ~ 4) which is roughly twice the theoretical uncertainty assumed in the current SZ power spectrum analysis (Reichardt et al. 2011). The first case is

where the pressure profile amplitude is maximally suppressed by high feedback and non-thermal pressure and low concentration: $\epsilon_f = 10^{-5}$, $\epsilon_{DM} = 0.1$, $\alpha_0 = 0.3$, and $A_c = 0.8$. The second case other case consists of the parameter choice that predicts a higher amplitude: $\epsilon_f = \epsilon_{DM} = \alpha_0 = 0.0$ and $A_c = 1.2$. We find that even for the two extreme gas physics scenarios the $A_{\text{SZ}} - B_{\text{tSZ}}$ has an overall uncertainty of only by $\sim 15\%$ (indicated by dashed lines in Fig 6). The slope of the relation changes from 1.33 to 1.47 between these two cases. This is encouraging as using the skewness spectrum data we can measure B_{tSZ} , then use this theoretical relation to derive A_{tSZ} .

The relation in Eq. 17 implies the skewness spectrum amplitude varies as σ_8^{11-12} . We also investigate how the skewness spectrum varies with other cosmological parameters. Figure 7 shows the variation of the skewness spectrum with σ_8 and Ω_b . Each panel shows how each cosmological parameter changes the skewness spectrum while other parameters are fixed at their respective fiducial values. The figure shows that the skewness spectrum depends very sensitively on σ_8 , and the amplitude varies as $B_{\text{SZ}}^\Lambda \propto \sigma_8^{11.6}$, as expected. The variations with Ω_b is moderate, while Ω_m , h , w_0 , and n_s show much smaller variation. We find the SZ amplitude for the skewness spectrum follows the scaling (at $\ell = 3000$),

$$B_{\text{tSZ}} \propto \left(\frac{\sigma_8}{0.8}\right)^{11.6} \left(\frac{\Omega_b}{0.045}\right)^{4.1} \left(\frac{h}{0.71}\right)^{2.5} \left(\frac{w_0}{-1.0}\right)^{-0.95} \times \left(\frac{n_s}{0.97}\right)^{-1.5} \left(\frac{\Omega_m}{0.27}\right)^{-0.46}. \quad (18)$$

Note that not all cosmological parameters follow a simple $B_{\text{tSZ}} = A_{\text{tSZ}}^{1.4}$ relation. The difference in scaling is due to the different range of masses in the mass function contributing to the skewness spectrum and the power spectrum. Note that we ignore the cosmology dependence of the concentration-mass relation (Bhattacharya et al. 2011a) in our study which can lead to a slightly different scaling of the skewness spectrum amplitude with cosmology. We provide the theoretical prediction of the skewness spectrum in Table 3 for our fiducial cosmology and gas parameters.

5. DETECTABILITY

We now assess the detectability of the SZ skewness spectrum in the presence of other secondary CMB anisotropy signals and extra-galactic foregrounds. Other than the thermal SZ effect, point sources (dusty, star-forming galaxies (DSFGs), and radio sources) will also contribute to the skewness spectrum signal. The main difference is while the thermal SZ effect contribute a negative bispectrum, the point sources contribute a positive signal (Rubio-Martín & Sunyaev 2003). Assuming that the clustered component of the DSFGs follows the Gaussian distribution, the skewness spectrum of the correlated DSFG is zero. As mentioned before, the kinetic SZ follow approximately Gaussian distribution and is thus assumed to have negligible skewness. We also ignore a possible skewness spectrum signal arising from the correlation between the SZ effect and the DSFG sources.

5.1. Noise Estimates

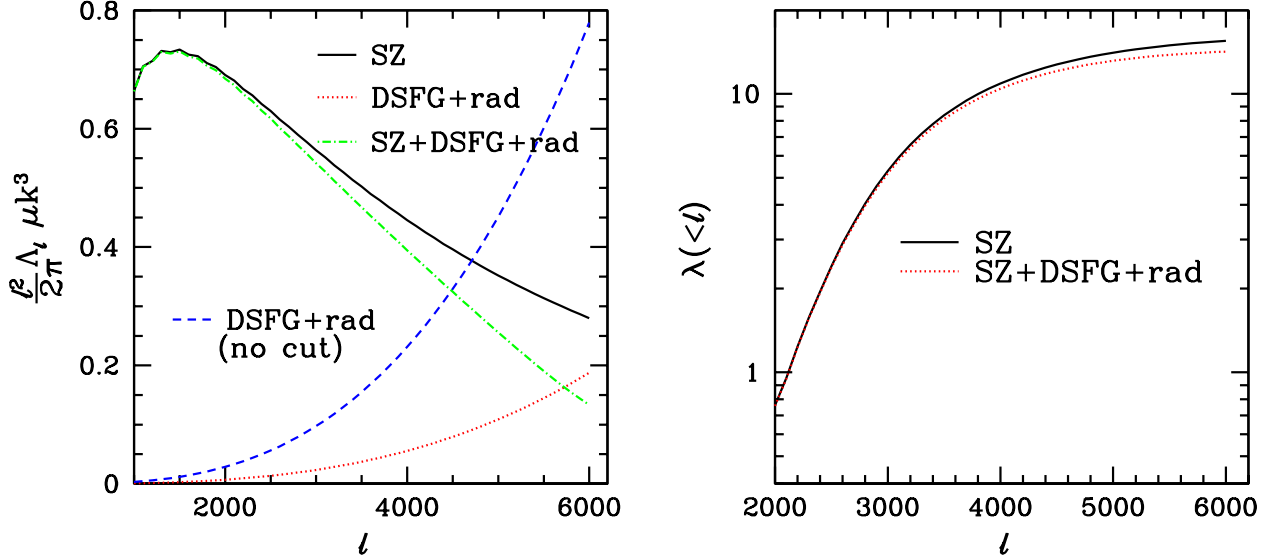


FIG. 8.— Detectability of the SZ skewness spectrum at 150 GHz. The left panel shows the SZ (black solid) and the point sources-DSFG+radio skewness spectrum (red dotted line). Also shown are the total skewness spectrum: SZ+DSFG+radio (green thin dot-dashed) and the DSFG+radio skewness spectrum if no point sources are removed from the map (blue short dashed line). The right panel shows the signal-to-noise integrated to a certain ℓ of the SZ skewness spectrum (black solid), the total- SZ+DSFG+radio (red dotted). We assume an SPT-like survey with sky coverage of 2500 deg^2 , $18 \mu\text{K-arcmin}$, 1.2 arcmin at 150 GHz.

The noise of the total skewness spectrum, assuming the primary contribution to the variance of the bispectrum comes from the power spectrum of the CMB sky, is given by Komatsu & Spergel (2001)

$$N^2(\ell_1 \ell_2 \ell_3) = C(\ell_1)C(\ell_2)C(\ell_3)\Delta_{\ell_1 \ell_2 \ell_3}, \quad (19)$$

where $C(\ell)$ is the total power spectrum that includes contributions from (1) the SZ power spectrum computed using Eq. 2, (2) the lensed CMB power spectrum computed using CAMB for the fiducial cosmology (Lewis et al. 2000), (3) the noise due to the finite resolution of an experiment, and (4) point sources, including dusty star-forming galaxies (DSFG) and radio sources. Note that $\Delta_{\ell_1 \ell_2 \ell_3} = 6$ if all three ℓ s are equal, $\Delta_{\ell_1 \ell_2 \ell_3} = 2$ if two of the three ℓ s are equal and $\Delta_{\ell_1 \ell_2 \ell_3} = 1$ if all three ℓ s are different. We describe the estimate of (3) and (4) in more detail below.

The noise due to the finite resolution of the experiment is given by Knox (1995); Jungman et al. (1996)

$$N_b(\ell) = w^{-1} \exp \left[\frac{\ell^2 \theta^2}{8 \log 2} \right], \quad (20)$$

where $w^{-1} = [\sigma_{\text{pix}} \theta / T_{\text{CMB}}]^2$ is the weight per unit solid angle, σ_{pix} is the noise per pixel, θ is the FWHM of the instrument in radians, T_{CMB} is the CMB temperature. Eq. 19 gives the expression for noise in the full sky limit. For the fractional coverage, Eq. 19 is multiplied by $f_{\text{sky}}^{-3/2}$, where f_{sky} is the fraction of the sky covered by the assumed surveys. Here we consider an SPT like survey with f_{sky} given by 2500 deg^2 sky coverage, σ_{pix} is $18 \mu\text{K-arcmin}$ at 150 GHz with the pixel size set by a Gaussian beam of FWHM=1.2 arcmin. The powers of the unresolved DSFG sources at 150 GHz are given by $5.25 \times 10^{-6} \mu\text{K}^2$ (Poisson) and $4.4 \times 10^{-6} (\ell/3000)^{-1.2} \mu\text{K}^2$ (clustered), and that of the radio sources is $9 \times 10^{-7} \mu\text{K}^2$

(Poisson) at 150 GHz (Reichardt et al. 2011).

5.2. Point Source Bispectrum

The DSFG population comprises of a clustered component and a Poisson component, both of which contribute to the point source power spectrum. We assume that the clustered component produces Gaussian fluctuations and hence does not contribute to the bispectrum of the point sources. We thus consider the bispectrum contribution only from the Poisson component of the DSFGs and the radio galaxies. The Poisson component of the DSFG typically comprises of two populations: (1) the bright point sources which are detected by the CMB experiments (Vieira et al. 2010; Marriage et al. 2011) and (2) the faint sources which follow a shallower distribution compared to the bright sources, and are undetected by the current CMB experiments but contribute a significant portion of the arcminute scale DSFG power measured in the CMB experiments. Both populations are expected to contribute to the total point source bispectrum.

Given the DSFG flux number counts dN/dS , we can write the bispectrum of DSFGs as

$$\frac{b_{\text{DSFG}}}{g^3(x_\nu)} = \int_0^{S_{\text{min}}} dS S^3 \frac{dN}{dS} = \frac{N_{\text{DSFG}}(> S_{\text{min}})}{3/\beta_{\text{DSFG}} - 1} S_{\text{min}}^3, \quad (21)$$

where the flux number counts is assumed to vary as $dN/dS \propto S^{-\beta-1}$, S_{min} is the flux cut above which all sources are detected and removed from the CMB maps, $N_{\text{DSFG}}(> S_{\text{min}})$ is the number density (per unit solid angle) of the DSFG with $S > S_{\text{min}}$, $g(x_\nu) = 1/(65.55 \text{ MJy sr}^{-1})[\sinh(x_\nu/2)/x_\nu^2]^2$, and $x_\nu = \nu/(56.8 \text{ GHz})$. Similarly, we define the bispectrum of the radio sources in terms of the radio source power spectrum

as

$$\frac{b_{\text{rad}}}{g^3(x_\nu)} = \int_0^{S_{\text{min}}} dS S^3 \frac{dN}{dS} = \frac{N_{\text{rad}}(> S_{\text{min}})}{3/\beta_{\text{rad}} - 1} S_{\text{min}}^3. \quad (22)$$

We assume all point sources above 5σ at 150 GHz are detected and removed (masked). For the SPT 150 GHz band, this corresponds to $S_{\text{min}}=6.4$ mJy (Vieira et al. 2010). We adopt the number density of the DSFG to be $N_{\text{DSFG}} = 0.3 \text{ deg}^{-2}$ (Negrello et al. 2007) and $\beta_{\text{DSFG}} = 1.93$ for the total (bright + faint) DSFG population to be consistent with the total DSFG power spectrum measured in the CMB surveys (e.g., Reichardt et al. 2011). For the radio population we adopt $N_{\text{rad}}(> S_{\text{min}}) = 1.29 \text{ deg}^{-2}$ and $\beta_{\text{rad}} = 1.03$, which is consistent with the 150 GHz source counts in Vieira et al. (2010). The skewness spectra of the radio sources and DSFGs are calculated the same way as the SZ skewness spectrum by plugging Eqs. 21 and 22 in Eq. 10.

The skewness spectrum for the various components at 150 GHz are shown in the left panel of Figure 8. Provided the point sources that are detected above 5σ in the SPT band-power are removed, the DSFG and radio population make negligible contribution to the skewness spectrum at $\ell \lesssim 4000$. Beyond this scale the skewness spectra of point sources and SZ effect become comparable. If no point sources are removed, then the skewness spectrum of the point sources (DSFG+radio) is a factor of 4-5 higher compared to the case with point source removal and becomes the dominant signal at $\ell > 5000$. Since the point sources and the SZ bispectrum have opposite signs, the total skewness spectrum (SZ+point sources) is smaller than the SZ skewness spectrum alone, especially at higher- ℓ where the point source bispectrum becomes non-negligible.

5.3. Signal-to-Noise

Finally, using Eq. 7, we define the signal-to-noise (S/N) in terms of the full-sky bispectrum. The S/N integrated to a certain ℓ (Komatsu & Spergel 2001; Hu 2000), $\lambda(< \ell)$ or λ_ℓ can be written as

$$\lambda(< \ell) = \sqrt{\sum_{\ell_1 \ell_2 \ell_3} \frac{B^2(\ell_1 \ell_2 \ell_3)}{N^2(\ell_1 \ell_2 \ell_3)}}, \quad (23)$$

where $\ell_3 \leq \ell_2 \leq \ell_1 \leq \ell$. The right panel of Figure 8 shows the total S/N for the SZ skewness spectrum. For a 2500 deg^2 SPT like survey at 150 GHz, the integrated S/N for the SZ skewness spectrum is ≈ 16 . Note that the total (SZ+point source) skewness spectrum is lower than the SZ skewness spectrum alone, because these two spectra have opposite sign. We find the S/N of the total skewness spectrum is only 10% smaller than the SZ only case.

6. SUMMARY AND DISCUSSIONS

We investigated the SZ bispectrum and the prospects of its detection with the current and upcoming CMB surveys. We define a skewness spectrum of the SZ effect which is a sum over all possible triangle configurations over the two smaller sides and expressed as a function of the largest side. About half of the SZ skewness spectrum signal arises from massive ($M_{500} \gtrsim 6.5 \times 10^{14} h^{-1} M_\odot$)

clusters in the local universe ($z \lesssim 0.4$). This is in contrast to the SZ power spectrum whose signals come primarily from high- z , low-mass galaxy groups ($M_{500} < 2 \times 10^{14} h^{-1} M_\odot$ and $z \gtrsim 0.6$) that have little observational constraints at present. The skewness spectrum, therefore, is less susceptible to the large astrophysical uncertainties associated with high- z , low-mass groups.

Our model for the SZ profile from groups and clusters reproduces the pressure profile of massive X-ray clusters in the local universe. This provides significant constraints on the ICM models of the present-day massive clusters leaving only the redshift dependent parameters unknown. Adopting wide priors on these parameters, we determine astrophysical uncertainties in the skewness spectrum to be about 33%. This is in contrast to the SZ power spectrum which is uncertain at the level of 50%, primarily because a significant fraction of the power comes from high- z , low-mass groups with $M_{500} \lesssim 2 \times 10^{14} h^{-1} M_\odot$ and $z \gtrsim 0.6$. Such objects have not been studied in as much detail as their high mass counterparts, making their theoretical modeling and hence the SZ power spectrum highly uncertain. Other than the gas physics uncertainties, the theoretical modeling of the cluster mass function also add to the uncertainties in the bispectrum prediction. Because of the exponential nature of the mass function at the massive cluster end, a small change in cluster masses cause significant difference in the abundance of massive clusters (Bhattacharya et al. 2011).

The relation of the amplitude of the SZ skewness spectrum and the SZ power spectrum is robust to these astrophysical uncertainties. Indeed, the overall uncertainty of the relation is only 7% within the allowed range of ICM parameters. The skewness and the power spectrum amplitudes of the thermal SZ effect is related as $B_{\text{SZ}}^\Lambda = A_{\text{tSZ}}^{1.4}$; consequently, the skewness spectrum varies as $B_{\text{SZ}}^\Lambda \propto \sigma_8^{11.6}$. Given the astrophysical uncertainty of 33%, the skewness spectrum measurements have the potential to constrain σ_8 to better than 3% accuracy. Since the SZ skewness spectrum gets its contribution primarily from the thermal component, the SZ skewness spectrum measurements can constrain the thermal SZ component uniquely. We can then use the skewness spectrum measurements to break the degeneracy between thermal and the kinetic power spectrum measurements.

We show that the current high resolution experiments, such as ACT, Planck, and SPT, have the potential to measure the skewness spectrum. For example, using the full 2500 deg^2 of data, SPT should be able to detect the SZ skewness spectrum with a signal-to-noise ≈ 16 . Indeed, the SZ skewness signal we predict in this study is detected by the ACT collaboration at 5σ (Wilson et al. 2012). We note however that our noise calculation does not account for all possible sources, as we make several simplified assumptions. For example, the contribution of the trispectrum is assumed to be negligible and we ignore any spatial correlation between the tSZ and CIB. However, as most of the SZ trispectrum signal comes from massive clusters (Shaw et al. 2009), masking out these objects should help reduce the SZ trispectrum contribution to the non-Gaussian component of the variance of the skewness spectrum. Also we assume the lensing 3-pt function and its correlation with the SZ 3-pt function

is zero, because CMB lensing dominates at $\ell \leq 2000$ while the SZ skewness spectrum gets most of the signal at $\ell \sim 4000$ (see Fig 8, right panel). In this study we also assume the kSZ bispectrum is negligible. While this assumption is valid for the linear part of the kSZ effect, the non-linear kSZ effect contributes a non-zero bispectrum which is not included in our calculation. However, given that the overall magnitude of the kSZ effect is about 20-40% of the thermal SZ amplitude, the non-linear kSZ bispectrum should be small. Nevertheless, in order to use the SZ skewness spectrum as a robust cosmological probe, these assumptions must be tested using mock catalogs. We will present such study in our future work.

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